

The effect of the fluorescence yield selection on the energy scales of Auger, HiRes and TA

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Abstract. The fluorescence yield data used for shower reconstruction in the Auger, HiRes and TA experiments are different, not only in the overall absolute value but also in the wavelength spectrum and the various atmospheric dependencies. The effect on the energy reconstruction of using different fluorescence yield parameterizations is discussed. In addition, the impact of a change in the fluorescence spectrum depends on the optical efficiency of the telescopes. A simple analytical procedure allows us to evaluate the combined effect of fluorescence yield and optical efficiency showing a non-negligible deviation between the energy scales of TA and Auger. However no relevant effect is found in the comparison between HiRes and Auger. Finally we show that a similar procedure could also be applied with real data.

1. INTRODUCTION

The accurate measurement of the calorimetric energy of UHECR showers by means of fluorescence telescopes relies on the precise knowledge of the air-fluorescence yield. In practice, the fluorescence yield is described by a set of parameters that allows us to convert the fluorescence intensity recorded by the telescope into energy deposited in the atmosphere at given atmospheric conditions. Auger, HiRes and Telescope Array (TA) use different parameterizations of the fluorescence yield (absolute value, wavelength spectrum and atmospheric dependency). The comparison of the energy spectra of Auger and HiRes/TA indicates a possible disagreement in the energy scale [1], which could be partly due to using different fluorescence yield data. In this work we evaluate the effect of the fluorescence yield choice on the relative energy scale of these experiments.

While deviations in the absolute value of the fluorescence yield translate directly into deviations in the reconstructed energy, the effect of differences in the relative wavelength spectra depends on the optical efficiency of the telescope. Therefore, in order to evaluate the actual influence of the fluorescence yield on the energy scales of these experiments, the combined effect of the assumed Y_λ values (i.e., absolute spectral values) and the optical efficiency ε_λ has to be analysed.

The dependence of the fluorescence yield with the atmospheric parameters is determined by the characteristic pressures P'_λ . Neglecting the humidity contribution to quenching and the temperature dependence of the collisional cross section, the differences in the assumed P'_λ values do not lead to significant disagreements in the energy reconstruction, since most of the shower energy is deposited in the lower atmospheric layers where quenching effects saturate. However these humidity and temperature contributions to P'_λ give rise to a non-negligible effect on the reconstructed energy [2].

The wavelength spectrum assumed by Auger, HiRes and TA, as well as the corresponding optical efficiencies of these experiments are described in section 2. The analytical method we used to compute the differences in the energy scales is described in section 3 and the corresponding results are presented in section 4. This theoretical method overestimate somewhat the energy deviations with respect to those

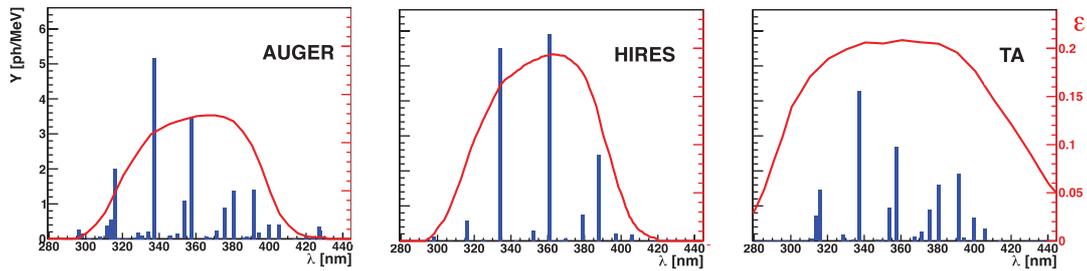


Figure 1. Absolute fluorescence yield (dry air) spectra at 800 hPa and 293 K employed by the Auger, HiRes and TA experiments (blue bars). The red line represents the optical efficiency of the corresponding telescopes including all optical elements.

from a full reconstruction of real showers. In section 5 we discuss on the applicability of this procedure to real data.

2. FLUORESCENCE YIELD DATASETS AND OPTICAL EFFICIENCIES

Data on the fluorescence yield used by Auger, HiRes and TA and the optical efficiency of the corresponding telescopes are described below. These parameters are shown together in Figure 1 for comparison.

2.1 Auger

The fluorescence yield model presently used by the Pierre Auger Collaboration is defined as follows. The absolute value of the yield for the 337 nm band Y_{337} is the one measured by Nagano et al. [3]. The yields for the remaining wavelengths are distributed according to the relative intensities measured by the AIRFLY Collaboration [4]. The characteristic pressures P'_λ are also those from AIRFLY [4]. Auger is the only experiment taking into account the water vapor quenching of the atmosphere and the temperature dependence of the collisional cross section [5].

The relative optical efficiency of the Auger telescopes [6] is dominated by the filter (MUG6 glass manufactured by Schott-Desag) although other elements (i.e., mirror, corrector ring, and PMT) also contribute.

2.2 HiRes–TA (Middle Drum)

The HiRes Collaboration uses the absolute value of the integrated fluorescence yield (300–400 nm) from Kakimoto *et al.* [7]. The relative intensities of the 337, 351 and 391 nm bands are those measured by Kakimoto while the intensities of the remaining spectral components are distributed according to those from Bunner [8]. The P'_λ values are those reported by Kakimoto [7].

The optical efficiency of the HiRes telescopes is described in [9]. The filter is very similar to the one used by Auger. One of the fluorescence detectors of HiRes has been refurbished and is presently used by the Telescope Array Collaboration, i.e., the so-called Middle Drum (MD) telescope.

2.3 Telescope Array (Black Rock Mesa & Long Ridge)

The Black Rock Mesa (BRM) and Long Ridge (LR) eyes of the TA Collaboration use the absolute value of the integrated fluorescence yield (300–400 nm) from Kakimoto et al. [7], that is, the same absolute reference of HiRes. However TA uses the wavelength spectrum measured by FLASH [10].

This spectrum has two additional bands not reported by Kakimoto that are added by TA to get their total absolute yield. The P'_λ values are also those from Kakimoto et al. [7].

These BRM and LR telescopes use a BG3 glass (manufactured by Schott) as UV filter. This filter has a higher transmittance over a wider wavelength interval than the MUG-6 glass employed by Auger. This is a relevant feature that potentially leads to a lower signal-to-noise ratio. In addition, other molecular bands, not taken into account because they are absorbed by other narrower UV filters, can contribute to the total fluorescence signal recorded by TA.

3. METHOD

A first estimate of the effect of a change in the fluorescence yield on the reconstructed energy can be obtained as follows. Neglecting other contributions (e.g., atmospheric transmission and geometrical factors), the fluorescence signal recorded by a telescope due to a *true* energy deposition E_t is proportional to $E_t \sum_\lambda Y_{\lambda t} \varepsilon_\lambda$, where ε_λ is the optical efficiency of the telescope and $Y_{\lambda t}$ is the *true* fluorescence yield, both dependent on wavelength λ . Hereafter we relax the notation using $\mathbf{Y}\boldsymbol{\varepsilon} \equiv \sum_\lambda Y_\lambda \varepsilon_\lambda$.

Let us assume two experiments with efficiencies $\boldsymbol{\varepsilon}_1$ and $\boldsymbol{\varepsilon}_2$ using different fluorescence yields \mathbf{Y}_1 and \mathbf{Y}_2 respectively. For a given E_t value, the first experiment reconstructs an energy E_1 . If the fluorescence yield is replaced by \mathbf{Y}_2 , the reconstructed energy would be E_{21} , that is,

$$E_t \mathbf{Y}_1 \boldsymbol{\varepsilon}_1 = E_1 \mathbf{Y}_1 \boldsymbol{\varepsilon}_1 = E_{21} \mathbf{Y}_2 \boldsymbol{\varepsilon}_1. \quad (1)$$

Therefore the effect of a change in the fluorescence yield is given by

$$\frac{E_{21}}{E_1} = \frac{\mathbf{Y}_1 \boldsymbol{\varepsilon}_1}{\mathbf{Y}_2 \boldsymbol{\varepsilon}_1}. \quad (2)$$

A relationship similar to (1) applies to experiment 2

$$E_t \mathbf{Y}_1 \boldsymbol{\varepsilon}_2 = E_2 \mathbf{Y}_2 \boldsymbol{\varepsilon}_2. \quad (3)$$

In principle the ratio of the energies measured by these two telescopes E_2/E_1 cannot be extracted from the above expressions. Nevertheless if the *true* fluorescence yield is known, expressions (1) and (3) would lead to the solution

$$\frac{E_2}{E_1} = \frac{\mathbf{Y}_1 \boldsymbol{\varepsilon}_1 \mathbf{Y}_1 \boldsymbol{\varepsilon}_2}{\mathbf{Y}_2 \boldsymbol{\varepsilon}_2 \mathbf{Y}_1 \boldsymbol{\varepsilon}_1}, \quad (4)$$

and thus, in case one of the two experiments is assumed to use the *true* fluorescence yield (e.g., $\mathbf{Y}_1 = \mathbf{Y}_t$), equation (4) will give the simple result

$$\frac{E_2}{E_1} = \frac{\mathbf{Y}_1 \boldsymbol{\varepsilon}_2}{\mathbf{Y}_2 \boldsymbol{\varepsilon}_2}. \quad (5)$$

This is similar to equation (2), although for the combined effect of efficiency and yield the ratio has to be computed with the efficiency of the experiment using the *wrong* yield. Note that even if the assumed *true* yield turns out to be wrong by a scale factor, expressions (4) and (5) are still valid, that is \mathbf{Y}_t only needs to describe the *true* relative fluorescence spectrum.

The above procedure allows the evaluation of the deviation in the energy deposition at a given atmospheric depth due to a change in the fluorescence yield (2) or in both fluorescence yield and optical efficiency (4, 5). The corresponding effect on the total calorimetric energy of the shower can also be calculated from the above expressions as long as we replace the above definition, $\mathbf{Y}\boldsymbol{\varepsilon} \equiv \sum_\lambda Y_\lambda \varepsilon_\lambda$, by the following one

$$\mathbf{Y}\boldsymbol{\varepsilon} \equiv \int_X dX \frac{dn}{dX} G(X) \sum_\lambda Y_\lambda(X) \varepsilon_\lambda T_\lambda(X), \quad (6)$$

Table 1. Energy deviation with respect to Auger ΔE , defined in (7), for different optical efficiencies (rows) and fluorescence yields (columns). Values (black) in the upper row represent the energy deviation if Auger would have used the fluorescence model of TA or HiRes. Assuming that Auger uses the *true* fluorescence spectrum, the energy scale of HiRes (TA) is shifted with respect to that of Auger by the blue (red) value. Results are shown for three cases: fixed pressure and temperature, realistic Gaisser-Hillas shower without and with temperature and humidity effects.

		800 hPa 293 K		10 ¹⁹ eV proton w/o T-h		10 ¹⁹ eV proton T-h	
		HiRes	TA	HiRes	TA	HiRes	TA
ε_λ	Y_λ						
Auger		5	12	3	9	~ 0	6
HiRes		5		3		~ 0	
TA			16		13		10

where dn/dX is the normalised development of energy deposited in the atmosphere by the shower, G is the geometrical factor, X is the slant depth and T_λ is the atmospheric transmission of the light traversing the path from the particular shower point to the telescope location. Using (6) has allowed us to calculate more realistically the energy deviations. For the development of deposited energy, a Gaisser-Hillas profile has been used and the T_λ values have been computed from [11] (see [2] for details).

4. RESULTS

In the first place we present the results neglecting Rayleigh scattering and geometrical factors. Using the data of section 2, we have obtained the $\mathbf{Y}\boldsymbol{\varepsilon}$ values for all the yield-efficiency combinations. The percentage energy deviations with respect to that of Auger, defined in this work as

$$\Delta E = 200 \times \frac{E - E^{\text{Auger}}}{E + E^{\text{Auger}}}, \quad (7)$$

have been evaluated at fixed atmospheric conditions (800 hPa and 293 K) for both HiRes and TA (see Table 1). The results in the first row have been obtained using the optical efficiency of Auger and therefore they represent the energy effect if the Auger yield were replaced by those of the other experiments (2). Those of second and third rows have been evaluated using the efficiencies of HiRes and TA respectively and therefore they give us the relative energy scale of these experiments with respect to Auger if we assume that Auger is using the *true* fluorescence spectrum (5).

Next a more realistic calculation has been carried out by averaging the energy ratios on the shower development including the Rayleigh effect and geometrical factors, that is, we have used the $\mathbf{Y}\boldsymbol{\varepsilon}$ definition given in expression (6). Calculations for a vertical 10¹⁹ eV proton shower as a function of the distance between the telescope and the shower maximum in the range 600–2600 g cm⁻² have been performed. The results for the above yield-efficiency combinations are plotted in figure 2. As shown in the figure, ΔE depends on the mass thickness between the telescope and the shower maximum due to the Rayleigh effect. The corresponding ΔE values for a typical vertical shower at 15 km from the telescope are also shown in the table (columns 3 and 4).

These calculations have been repeated including the temperature and humidity contributions to the P'_λ values in Auger with typical atmospheric profiles (see [2] for details). As shown in the table (columns 5 and 6), the disagreement between the relative energy scales for both cases HiRes-Auger and TA-Auger is partly cancelled. This is expected since the reconstructed energy increases when the temperature and humidity is included.

In summary, assuming that the AIRFLY fluorescence spectrum is the *true* one, according to our calculations the relative energy scale of TA is 13% (10%) larger than that of Auger if the temperature and humidity effects are (not) included in the Auger yield. On the other hand, the Auger energy would

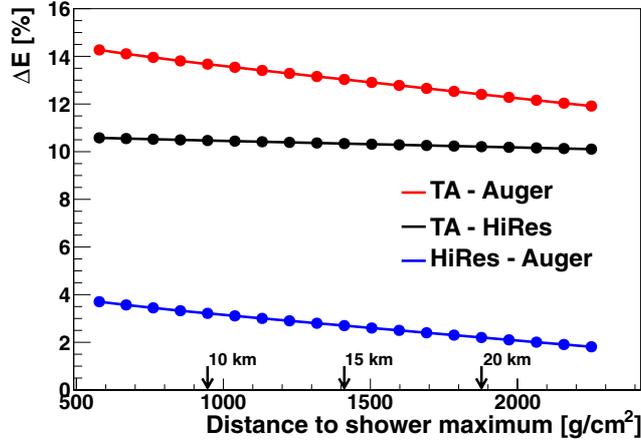


Figure 2. Relative deviations of the energy scales as a function of the distance between the telescope and the shower maximum for a vertical proton shower of 10^{19} eV. Rayleigh scattering has been included in this calculation.

increase by 10% (6%) if this experiment replaced its yield by that of TA, without (with) including temperature and humidity effects. Finally we infer that the energy scale of TA is shifted by about 10% with respect to that of HiRes.

5. TOWARDS A PRECISE DETERMINATION OF THE RELATIVE ENERGY SCALES

The above results on the relative energy scales of Auger, HiRes and TA have been obtained assuming ideal showers that only emit UV radiation from atmospheric fluorescence. In practice a non-negligible amount of Cherenkov light is emitted and partly detected by the telescope. The contribution of the Cherenkov light has to be included in the reconstruction algorithm to measure the shower energy [12]. The larger the Cherenkov contribution, the smaller the effect of a change in the fluorescence yield on the reconstructed calorimetric energy. As a consequence, a given change in the fluorescence yield does not translate entirely in the energy scale and therefore the above results are expected to overestimate the energy effect in real data.

The effect of a change in the fluorescence yield can easily be obtained from the observed deviation in the reconstructed energy of real showers when replacing in the reconstruction algorithm the Y_λ data by that of other experiment. However, the evaluation of the combined effect of optical efficiency and fluorescence yield is not trivial since replacing the optical efficiency in the reconstruction algorithm by that of a different experiment gives rise to a wrong reconstructed energy. However, as shown next, the problem can be solved following arguments similar to those presented in section 3.

The integrated signals (in appropriate units) recorded by experiments 1 and 2 from a given energy deposition in the atmosphere will be given by

$$S_1 = E_t Y_t \epsilon_1 = E_1 Y_1 \epsilon_1; \quad S_2 = E_t Y_t \epsilon_2 = E_2 Y_2 \epsilon_2. \quad (8)$$

In this case $Y\epsilon$ includes not only the fluorescence contribution given by equation (6) but also that of the Cherenkov light. In order to determine the relative energy scale of these experiments (E_2/E_1), the energy of given showers should be reconstructed by each experiment using the respective signal. Nevertheless, if we assume again that experiment 1 is using the *true* fluorescence yield, the relationship $S_2 = E_1 Y_1 \epsilon_2$ holds, that is, we could even calculate the signal of telescope 2 from that of telescope 1. Instead of that, the E_2/E_1 ratio can be obtained directly from expression (5) which also applies in this case. While the $Y\epsilon$ quantities were previously computed analytically for modelled showers (sect. 3), now we can

make the reconstruction algorithm to do it for us using real data, including in this way the Cherenkov contribution.

In general the reconstructed energy is proportional to the reciprocal $\mathbf{Y}\boldsymbol{\varepsilon}$ value, the telescope signal being the proportionality constant, i.e., $E = S/\mathbf{Y}\boldsymbol{\varepsilon}$. From the S_1 signal, E^* parameters can be obtained by reconstructing the energy assuming in the algorithm the optical efficiency of experiment 2 either with the fluorescence yield of experiment 1 or that of experiment 2,

$$E_1^* = \frac{S_1}{\mathbf{Y}_1} \boldsymbol{\varepsilon}_2; \quad E_2^* = \frac{S_1}{\mathbf{Y}_2} \boldsymbol{\varepsilon}_2. \quad (9)$$

Note that E_1^* and E_2^* are not actual values of energy deposition but those of E_1 and E_2 scaled by the common S_1/S_2 factor,

$$E_1^* = \frac{S_1}{S_2} E_1; \quad E_2^* = \frac{S_1}{S_2} E_2 \quad (10)$$

and therefore the relative energy scale can be determined from

$$\frac{E_2}{E_1} = \frac{E_2^*}{E_1^*}. \quad (11)$$

From arguments similar to those of section (3), the condition for experiment 1 of using the *true* fluorescence yield can be relaxed to just using the *true* fluorescence spectrum, although in this case we have to assume that the Cherenkov contribution is small with respect to that of fluorescence.

6. CONCLUSIONS

In this work we have evaluated the effect of the fluorescence yield choice on the energy reconstruction of extensive air showers. Using a simple analytical procedure we found that Auger (without T and h effects) would increase the energy by about 9% and 3% if the fluorescence yield of TA and HiRes respectively were used instead.

For the comparison of the energy scales, the optical efficiency of the experiments also matters. A simple procedure allows to carry out this comparison analytically. Assuming that Auger uses the *true* fluorescence spectrum, i.e., AIRFLY [4], we found that the TA energy is larger than that of Auger (including temperature and humidity effects) by about 10%. Notice that this energy deviation is accidentally lowered because TA does not take into account the temperature and humidity contributions to the fluorescence yield. Otherwise the energy scale of TA would be 13% larger.

The above results are somewhat overestimated because the analytical procedure shown here does not take into account the contribution of Cherenkov light in the reconstruction of real data. A procedure has been described that allows to make these comparisons from the reconstruction of real data of any experiment. Some tests have shown that the analytical ΔE values are overestimated with respect to those of real data by about 15% and therefore the energy scale of TA would be larger than that of Auger by about 6% (10%) if the humidity and temperature effects are (not) taken in account in Auger.

Finally, from the present analysis we also conclude that the energy of TA (BRM and LR eyes) is about 10% larger than that of HiRes (i.e., MD eye of TA) due to the different fluorescence spectra used by these experiments.

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