Analysis of atmospheric attenuation using the Telescope Array central laser data

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Abstract. Located in the western desert of the state of Utah, the Telescope Array (TA) experiment measures the properties of ultra high energy cosmic ray (UHECR) induced extensive air showers. TA employs a hybrid detector comprised of a large surface array of scintillator detectors overlooked by three fluorescence telescopes stations. The TA Low Energy extension (TALE) detector has operated as a monocular Cherenkov/fluorescence detector for nearly five years, and has recently been complemented by a closely spaced surface array to operate in hybrid mode. The TAx4 upgrade is underway and aims to, as the name suggests, quadruple the size of the surface array to improve statistics at the highest energies (post-GZK events).

The analysis of the TA fluorescence detectors (FD) data requires knowledge of the degree of the atmospheric attenuation of UV light produced by shower particles. This attenuation depends partially on the amount of aerosols present in the atmosphere at the time of shower observation. Being highly variable, real time measurement of the aerosols light attenuation is accomplished through the use of a central laser facility (CLF) located at the center of the surface array, and in the field of view of the three FDs, as well as, the TALE FD.

In this proceeding we will describe the experiment, and the CLF data and analysis, and give results on measured aerosols attenuation, yearly averaged. FD measurements of shower energy and $X_{\text{max}}$ involve corrections for atmospheric attenuation due to the presence of aerosols. We discuss the errors introduced into the shower parameters reconstruction due to uncertainty about aerosols attenuation.

1 Introduction

The Telescope Array experiment is located in the West Desert of Utah, about 150 miles southwest of Salt Lake City, and is the largest cosmic ray detector in the northern hemisphere. TA has operated since 2008. The detector consists of 507 surface detectors (SD), arranged in a square grid of 1.20 km spacing [1]. Three FD stations, comprised of 38 telescopes, are located at the periphery of the SD array [2, 3]. The FD telescopes observe the airspace above the SD array. This arrangement of detectors is shown in Figure 1. The goal of TA is to clarify the origin of ultra-high energy cosmic rays (UHECR) and related extremely high energy phenomena in the universe. The results of measuring the energy spectrum, composition, and anisotropy in the arrival direction distribution for energies above $10^{18.2}$ eV have been published [4–6].

At the time of this writing, TA is being expanded to an area roughly four times the size of the current array; this updated detector is referred to as TAx4. Some 500 surface detectors, scintillation counters, and 12 telescopes (two FD stations) are being constructed to accomplish the expansion.

2 Aerosols measurements using CLF

The TA experiment employs a laser system located at the center of the array to provide data for aerosols characterization. The laser points vertically and fires in a fixed direction. Typical laser pulse energy is a few milli-Joules,
bright enough to be seen by the three FD stations at a distance of ~21 km. An event recorded by the MD/TALE telescopes is shown in Figure 2 as an example.

An analysis on the CLF data observed by three FD stations over a number of years was carried out to measure the mean Vertical Aerosols Optical Depth (VAOD). CLF data is collected every half hour (300 laser shots). Each set of shots observed by one of the three FDs provides one measurement point. Data from each FD was averaged for each year of operation to obtain a mean VAOD as measured by that station. Results from the BR and LR stations are summarized in Table 1. Data from the MD station was also analyzed and compared to the other two stations. The results from MD were found to be consistent with those from the other two stations.

### 3 Shower Energy and $X_{\text{max}}$

The simulation and reconstruction of extensive air showers observed by the FDs employ an aerosols model which describes the attenuation of light using a vertical aerosols density and a horizontal extinction length. The aerosols density is assumed constant from ground level up to a “mixing-layer height”, $h_m$, decaying exponentially above $h_m$, with a scale height of $h_0$. The horizontal extinction length, $L$, gives the extinction length of light at a wavelength of 334 nm. A slowly varying function of wave-length corrects for different light wavelengths. The relation $VAOD = (h_m + h_0)/L$, connects the vertical optical depth to the model parameters. In practice, we set $h_m = 0$, and allow the other two parameters to vary. The choice of $L = 25$ km and $h_0 = 1$ km produces a mean $VAOD = 0.04$ and is the standard value used for various TA analyses. This value is consistent with those measured using the CLF data reported in section 2.

In this section we use MD observed data to examine the effect on the reconstruction of shower energy and $X_{\text{max}}$ of using this mean value of VAOD. Given the RMS values in Table 1, we compare the results to those obtained with VAOD values of 0.02 or 0.06. In addition we report on a MC study of using a mean value of VAOD in place of a time dependent measurement of aerosols. In particular, we simulate showers using aerosols distributions sampled from a database of actual measurements, and then reconstruct these showers using either the true model parameters or the mean value and compare the results.
For the energy and $X_{\text{max}}$ systematics check, we look at real shower data measured using the Middle Drum FD site in hybrid mode with the TA SD. Hybrid measurement insures accurate geometrical reconstruction, needed for reliable energy and $X_{\text{max}}$ reconstruction. This data sample, with additional quality cuts, was used as the basis of two TA publications [5, 7].

Keeping the aerosols model scale height at 1 km, we change the horizontal extinction length to produce a VAOD value of 0.02 or 0.06. We compare the reconstructed shower energy and $X_{\text{max}}$ on an event by event basis. Results are shown in Figure 3. As can be seen from the figure, we can quote a systematic uncertainty of roughly $\Delta E = 8.5\%$, and $\Delta X_{\text{max}} = 10 \text{ g/cm}^2$ due to aerosols.

The MD detector response to showers generated with a random sample of aerosols parameters drawn from the DB was performed. Nights where the aerosols concentration was too large (shower data would be discarded) were not included in the simulation; resulting in a mean value of 0.04 for the VAOD.

Detector observed events were then reconstructed using the mean VAOD value, $L = 25$ km and $h_0 = 1$ km, or using $L = L_{\text{thrown}}$ and $h_0 = h_0_{\text{thrown}}$. In both cases the thrown shower geometry was used in the reconstruction to mimic accurate geometrical reconstruction. A comparison of the two reconstructions is shown in Figure 5. As can be seen in the figure, the contribution to the reconstruction resolution due to the use of a mean VAOD can be estimated at $\Delta E = 12\%$, and $\Delta X_{\text{max}} = 13 \text{ g/cm}^2$.

### 4 Summary

A laser system located at the center of the TA detector (CLF) is used to measure the aerosols light attenuation. An analysis of the CLF data was performed to measure a yearly mean value for the VAOD. Results shown in Table 1. While variable in time, a mean value, $VAOD = 0.04$, is used in the analysis of shower data recorded by the TA FDs. A simple analysis of the effect of using this mean value on shower energy and $X_{\text{max}}$ gives estimates for systematic errors of roughly $\Delta E = 8.5\%$, and $\Delta X_{\text{max}} = 10 \text{ g/cm}^2$. Reconstruction resolution is affected at the levels of $\Delta E = 12\%$, and $\Delta X_{\text{max}} = 13 \text{ g/cm}^2$.

### References
